

Pre-Inflation in the Presence of Conformal Coupling

Apostolos Kuiroukidis*

and

Demetrios B. Papadopoulos

Department of Physics,

Section of Astrophysics, Astronomy and Mechanics,

Aristotle University of Thessaloniki,

541 24 Thessaloniki, GREECE

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Abstract

We consider a massless scalar field, conformally coupled to the Ricci scalar curvature, in the pre-inflation era of a closed FLRW Universe. The scalar field potential can be of the form of the Coleman-Weinberg one-loop potential, which is flat at the origin and drives the inflationary evolution. For positive values of the conformal parameter ξ , less than the critical value $\xi_c = (1/6)$, the model admits exact solutions with non-zero minimum scale factor and zero initial Hubble parameter. Thus these solutions can be matched smoothly to the so called Pre-Big-Bang models. At the end of this pre-inflation era one can match inflationary solutions by specifying the form of the potential and the whole solution is of the class $C^{(1)}$.

*E-mail: kouirouki@astro.auth.gr

I. Introduction

Inflation was historically introduced to solve the problem of monopoles that could be largely produced in the early Universe. It was subsequently elaborated into other forms such as the *chaotic inflation* scenario [1-3]. Inflation consists of a short period of accelerated superluminal expansion of the early Universe, at the end of which the transition to the standard big bang model should occur [4]. This apparently solves the problem of the flatness and smoothness of the Universe over such a large scale of distances [5].

In the scalar field driven inflationary scenario, the matter content of the Universe has the equation of state of the quantum vacuum, $P = -\rho$ (with P, ρ the pressure and energy density of matter). The evolution is regarded as the "rolling" of the value of the field which is minimally coupled to gravity, in the presence of the scalar field potential. The form of the potential is motivated usually from particle physics arguments. This is effectively a period of supercooling of the Universe where in fact through a phase transition the field rolls to its true vacuum state [3,5].

Usually the inflating Universe is taken to have flat spatial sections [3]. However recent WMAP data suggest a Universe with $\Omega_{tot} = 1.02 \pm 0.02$ [7]. This implies a closed ($k = +1$) model. Moreover using Penrose-Hawking-Geroch techniques one can evade certain theorems (see [11] and references therein) and allow for past-eternally inflating cosmological models that are also singularity free.

Indeed the problem of the avoidance of the initial singularity and of the exit from inflation are central to all the attempts to construct realistic inflationary models [11]. One of the advantages of the finiteness of the curvature invariant at the origin and of the zero Hubble parameter, is that these solutions can be matched smoothly to the so called Pre-Big-Bang models (see [9] and references therein).

In this paper we consider a positively curved FLRW metric, in the presence of a scalar field. The scalar field is conformally coupled to the spacetime scalar curvature invariant. We are mainly interested in the pre-inflation evolution era. This begins just after the Planck era ($T(GeV) \leq 10^{19}$) and ends at the supposed energy scale that the second order phase transition takes place (e.g., at the GUTs scale $T(GeV) \simeq 2 \times 10^{15}$). Then through the scalar field potential, the field rolls towards non-zero values, signifying the onset of inflation. In the pre-inflation era we assume that the *classical* value of the ϕ -field remains null, at the origin, ($\phi_c = 0$) and at the top of the potential profile $V(0)$. The specific form of the potential is not of particular importance here, since it occurs from various particle physics theories, as long as it is flat enough to drive inflation.

Now we assume that, although the scale factor in this pre-inflation evolution is a classical object, one must consider instead the time evolution of the quantum fluctuations of the ϕ -field instead. They can be large near the Planck time and as the scale factor increases we have the transition of the Universe towards clearly *classical* evolution, so these must be suppressed. This is a reasonable assumption [8]. The set of coupled equations for the scale factor and the time evolution of the

quantum fluctuations, that are suppressed as time evolves, is studied analytically and numerically. There exist families of solutions where the scale factor has a minimum non-zero value at the origin and the Hubble parameter is zero. These are particularly interesting because one does not enter in fact the quantum gravity regime at all [11]. Moreover they can be matched to the so called Pre-Big-Bang models smoothly [9]. Finally at the end of this stage of pre-inflation evolution, the Universe is at an inflationary state (i.e., the second derivative of the scale factor is non-negative ($R''(t) \geq 0$)) and one can match these solutions to classes of inflating solutions, specifying the form of the potential.

II. Inflation in FLRW Models

We consider a Universe described by the metric

$$ds^2 = -dt^2 + R^2(t) \left[\frac{dr^2}{1 - kr^2} + r^2(d\theta^2 + \sin^2\theta d\phi^2) \right], \quad (1)$$

of the Friedman-Lemaitre-Robertson-Walker (FLRW) class, where $R(t)$ is the scale factor. We use units where $\hbar = c = k_B = 1$ and set $\chi^2 := 8\pi G/3$ with $M_{Pl}^2 = G^{-1} \simeq (1.22 \times 10^{19} GeV)^2$ [3].

Also we assume the presence of a scalar field ϕ with energy density and pressure given by

$$\rho_\phi = \frac{\dot{\phi}^2}{2} + V(\phi), \quad P_\phi = \frac{\dot{\phi}^2}{2} - V(\phi), \quad (2)$$

where $V(\phi)$ is the potential. In the pre-inflation evolution the ϕ -field time evolution is extremely slow. The potential at the origin (i.e. for $\phi = 0$) is very flat, as for example in the Coleman-Weinberg (CW) one-loop, zero temperature case

$$V(\phi) = V(0) + B\phi^4 \left[\ln \left(\frac{\phi^2}{\sigma^2} \right) - \frac{1}{2} \right]. \quad (3)$$

Here $V(0) = B\sigma^4/2$ with $B \simeq 10^{-3}$ and $\sigma \simeq 2 \times 10^{15} GeV$ the Grand Unification scale. Also $\chi^2 V(0) \simeq (10^{10} GeV)^2$ [3].

In the pre-inflation era where $10^{15} \leq T(GeV) \leq 10^{19}$, gravity can be treated classically, but the other fields must be treated quantum-mechanically [8]. Here we employ the assumption that the ϕ -field is a conformally coupled massless scalar field in the action (see p. 44 of [6])

$$S = \int d^4x \sqrt{-g} \left[\mathcal{R} + 6\chi^2 \left(\frac{1}{2} g^{\mu\nu} \phi_{,\mu} \phi_{,\nu} - \frac{1}{2} \xi \mathcal{R}(t) \phi^2 - V(0) \right) \right]. \quad (4)$$

It obeys

$$[\square + \xi \mathcal{R}(t)] \phi(x) = 0. \quad (5)$$

Here $\mathcal{R}(t) = (6k/R^2(t))$ is the Ricci scalar curvature of the metric in Eq. (1). Usually, in front of the first term in the action functional, the factor $(1/16\pi G) = (1/6\chi^2)$ appears, which is set to unity by a proper choice of units. However we want the explicit Placnk-scale dependence to be present, so the precise form of the action functional should be the one above. The length scale of our model is simply $L_{Pl} = t_{Pl} = \chi$. The case $\xi = 0$ corresponds to the minimally coupled scalar field, while $\xi_c = (1/6)$ to the conformally coupled one.

The Einstein's equations are exactly given by (see p. 87 of [6])

$$G_{\mu\nu} + 3\chi^2 \left((1 - 2\xi)\phi_{;\mu}\phi_{;\nu} + (2\xi - \frac{1}{2}) g_{\mu\nu} g^{\rho\sigma} \phi_{;\rho}\phi_{;\sigma} - 2\xi\phi_{;\mu\nu}\phi + 2\xi g_{\mu\nu}\phi\Box\phi - \xi G_{\mu\nu}\phi^2 - g_{\mu\nu}V(0) \right) = 3\chi^2 T_{\mu\nu}. \quad (6)$$

The matter content of our model is given by $T_\nu^\mu = \text{diag}(\rho, -P, -P, -P)$ with $P = (\gamma - 1)\rho$. The flattness of the potential at the origin effectivelly appears like a cosmological constant.

The equations of motion are (for $k = +1$)

$$\ddot{\phi} + 3H\dot{\phi} + \frac{6\xi}{R^2(t)}\phi = 0, \quad (7)$$

and

$$H^2 + \frac{1}{R^2} = \frac{\chi^2}{(1 - 3\xi\chi^2\phi^2)} \left[\rho + \frac{\dot{\phi}^2}{2} + 6\xi H\phi\dot{\phi} + V(0) \right], \quad (8)$$

$$2\frac{\ddot{R}}{R} + H^2 + \frac{1}{R^2} = \frac{3\chi^2}{(1 - 3\xi\chi^2\phi^2)} \left[P + \left(\frac{1}{2} - 2\xi\right)\frac{\dot{\phi}^2}{2} + 4\xi H\phi\dot{\phi} - V(0) \right]. \quad (9)$$

Since it is generally accepted that in this era radiation dominates, we set $\gamma = (4/3)$. Eliminating ρ from Eqs. (8), (9) we obtain

$$2\frac{\ddot{R}}{R} = \frac{\chi^2}{(1 - 3\xi\chi^2\phi^2)} [(1 - 6\xi)\dot{\phi}^2 + 6\xi H\phi\dot{\phi} - 4V(0)]. \quad (10)$$

We will therefore consider in what follows, Eqs. (7) and (10). For a recent interesting discussion the reader is referred to [13].

III. Classes of Solutions from Numerical Integration

We introduce the parameters σ, \mathcal{E} through $V(0) = \sigma^4$ and $\chi^2 V(0) = \mathcal{E}^2$. For example in the case of the CW potential, σ is close to the Grand Unification scale and $\mathcal{E} = 10^{10} GeV$. We normalize the scale factor as $(R/R_{min}) := S \geq 1$ because we want a minimum scale factor and also define $(t/t_{Pl}) := \tilde{t} \geq 1$ because we assume that the

evolution takes place past the Planck energy scale. From similar considerations the combination $F := \chi\phi$ is dimensionless. Then Eqs. (7) and (10) give

$$F'' + 3\frac{S'}{S}F' + \frac{6\xi}{\lambda^2 S^2}F = 0, \quad (11)$$

$$2\frac{S''}{S} = \frac{1}{(1 - 3\xi F^2)} \left[(1 - 6\xi)(F')^2 + 6\xi\frac{S'}{S}FF' - \chi^2\mathcal{E}^2 \right], \quad (12)$$

where $\lambda := (R_{min}/\chi)$ and the prime denotes differentiation with respect to the \tilde{t} . Also $\delta := \chi\mathcal{E}$ is a dimensionless parameter (for example in the case of the CW potential

$$\delta := \chi\mathcal{E} = \frac{10^{10}GeV}{1.22 \times 10^{19}GeV} = 0.82 \times 10^{-9}). \quad (13)$$

So this parameter, (i.e., δ) which actually stems from the position of the symmetry breaking energy scale with respect to the Planck scale, is the first free parameter of our model. It gives the minimum scale factor in units of the Planck length. The other is of course $\xi \in [0, \frac{1}{6}]$.

Finally rescaling $T := \delta\tilde{t}$ and $\bar{\lambda} := (\lambda\delta)$ we arrive at a completely dimensionless form which is also esthetically pleasing

$$F'' + 3\frac{S'}{S}F' + \frac{6\xi}{\bar{\lambda}^2 S^2}F = 0, \quad (14)$$

$$2\frac{S''}{S} = \frac{1}{(1 - 3\xi F^2)} \left[(1 - 6\xi)(F')^2 + 6\xi\frac{S'}{S}FF' - 1 \right]. \quad (15)$$

We numerically integrate Eqs. (14) and (15), using as $\delta \simeq 10^{-9}$. We choose for convenience $\bar{\lambda} = 1$ which is the third free parameter of our model. The normalized time begins from $T_{init} = 10^{-9} \simeq 0$ i.e., just after the Planck time scale, up to $T_{fin} \simeq 0.1$ which is equivalent for the physical time scale t , to the onset of the Grand Unified era. We consider the time evolution of F as the suppression of the quantum fluctuations of the ϕ -field, around its zero classical value ($\phi_c = 0$), as time evolves. So the initial value should be $F_0 \simeq 1$ and the final $F_f \simeq 0$. Then we investigate the initial value of the field derivative $F'_0 = F'_0(\xi)$ as a function of ξ in order that the previous demand is fulfilled. Using double precision numerical integration codes, we find and plot the Scale factor $S(T)$ also. For $\tilde{\xi} = 0.05$ we have the maximum, in absolute value, field derivative $F'_0 = -12.8$. The time evolution of the ϕ -field and the scale factor for this particular value $\tilde{\xi} = 0.05$ are shown in Fig. 1 and Fig. 2. We stress again the fact that we seek solutions with non-zero minimum scale factor and zero initial Hubble parameter.

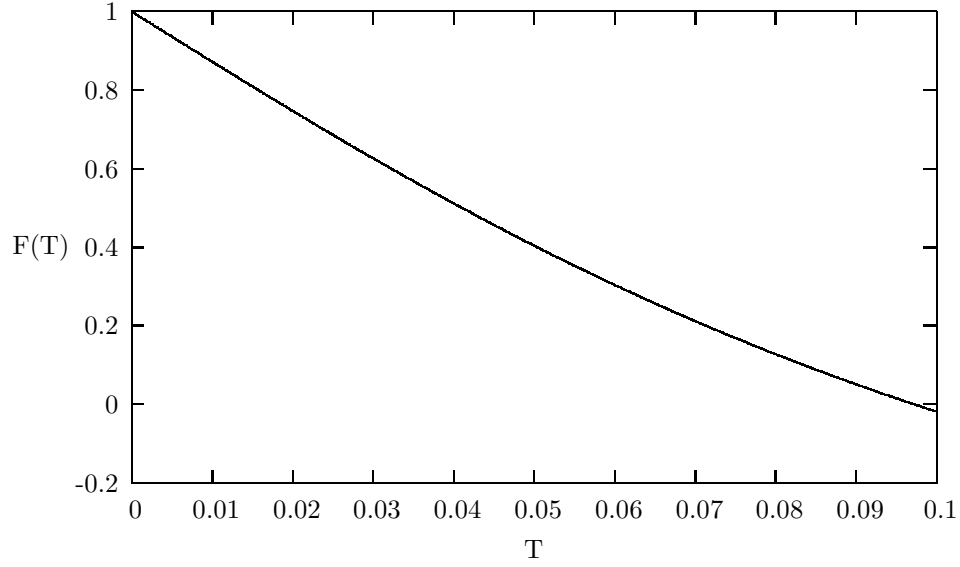


Figure 1: The normalized value of the quantum fluctuations, versus the normalized time T , of the field $F(T)$ for $\tilde{\xi} = 0.05$. They evolve towards zero value as the scale factor of the Universe increases.

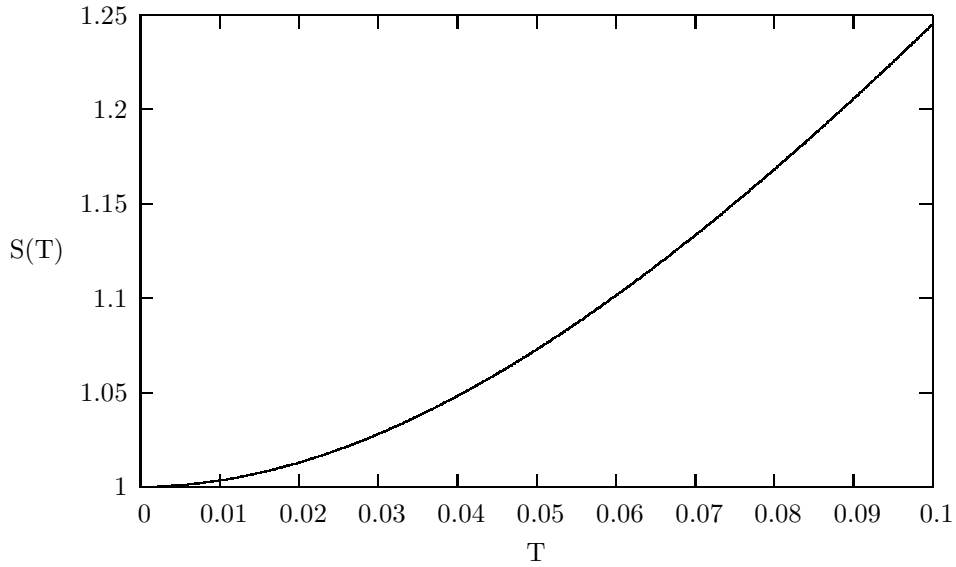


Figure 2: The normalized Scale factor $S(T)$, versus the normalized time T , for $\tilde{\xi} = 0.05$. The scale factor assumes a minimum value, with zero initial Hubble parameter.

Of course, every unit of the normalized time T corresponds to $\delta^{-1}t_{Pl} \simeq 10^9 t_{Pl}$ units of the physical time t . We emphasize that this depends on the relative value $V(0)$ with respect to the Planck energy scale.

In Fig. 3. the initial value of the field derivative F'_0 versus ξ is plotted for values $0 < \xi < (1/6)$. The numerical integration gives a rather accurate dependence as $F'_0(\xi) = 889(\xi - 0.05)^2 - 12.8$.

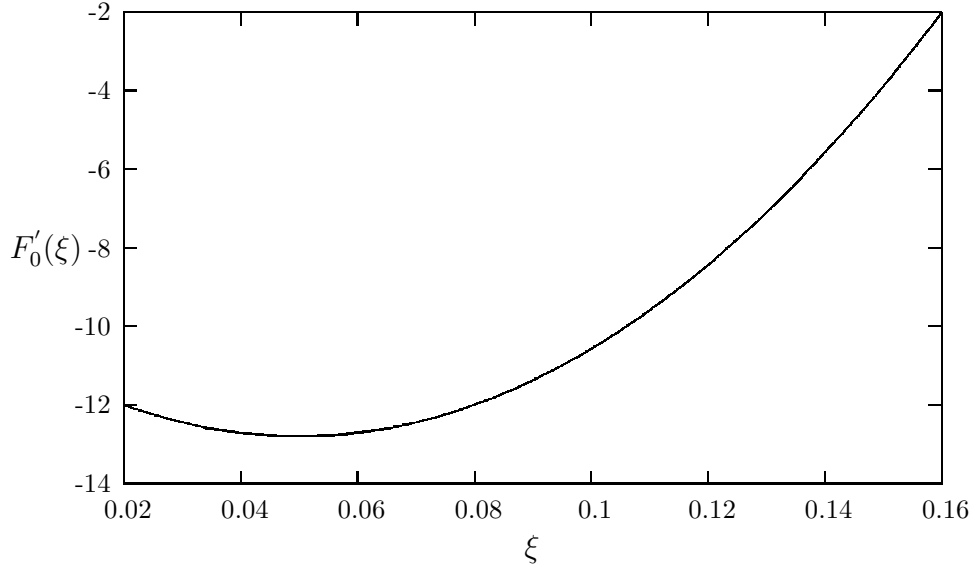


Figure 3: The initial time derivative of the field $F'_0 = F'_0(\xi)$.

In Fig. 4. the value of the normalized scale factor S_{fin} at the time $T \simeq 0.1$, versus ξ is plotted for values $0 < \xi < (1/6)$. This is effectively the normalized scale factor at the time the infationary era starts. One can therefore match here inflationary solutions with specific choices of the potential. The numerical integration gives a rather accurate dependence, through fitting using Mathematica[®], as $S_{fin}(\xi) = 1.344 - 2.232\xi + 0.8\xi^2$. As it is necessary $S_{fin} > 1$.

Also the value of the scale factor derivative $S'_{fin} = S'_{fin}(\xi)$ at the end is plotted in Fig. 5. It has the dependence as $S'_{fin}(\xi) = -38.8\xi + 6.043$. The value of the second derivative $S''_{fin} = S''_{fin}(\xi)$, as well as of the field time-derivative at the end of this period, $F'_{fin} = F'_{fin}(\xi) \simeq 0$, was found to be zero, independent of ξ , in the limits of the numerical accuracy achieved, at the end of the T -interval. So we are exactly at the end of an inflating process and here one can match inflationary solutions, with a given value of the Scale factor and the Hubble parameter, by choosing the appropriate value of ξ .

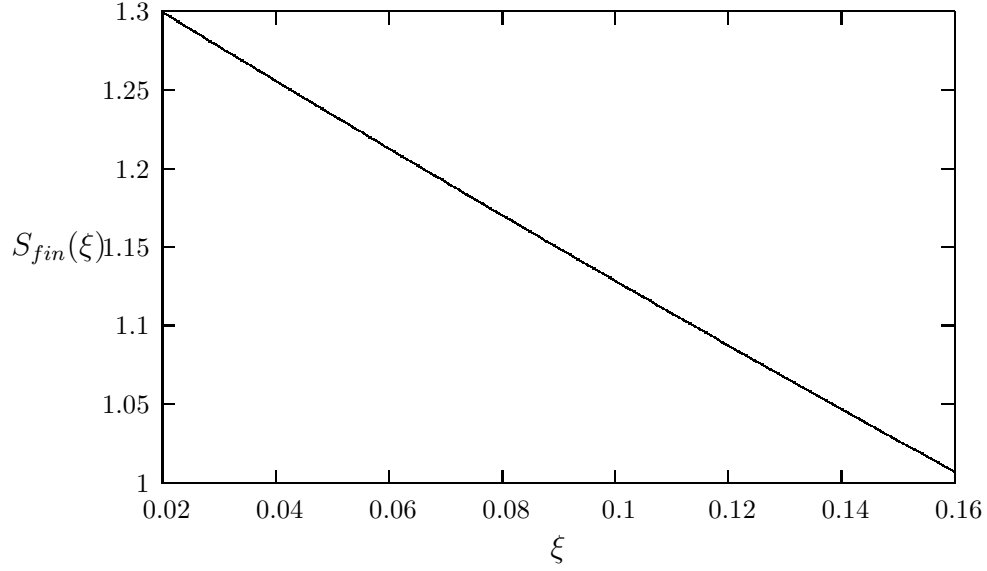


Figure 4: The Dependence $S_{fin} = S_{fin}(\xi)$ of the scale factor at the end of this period.

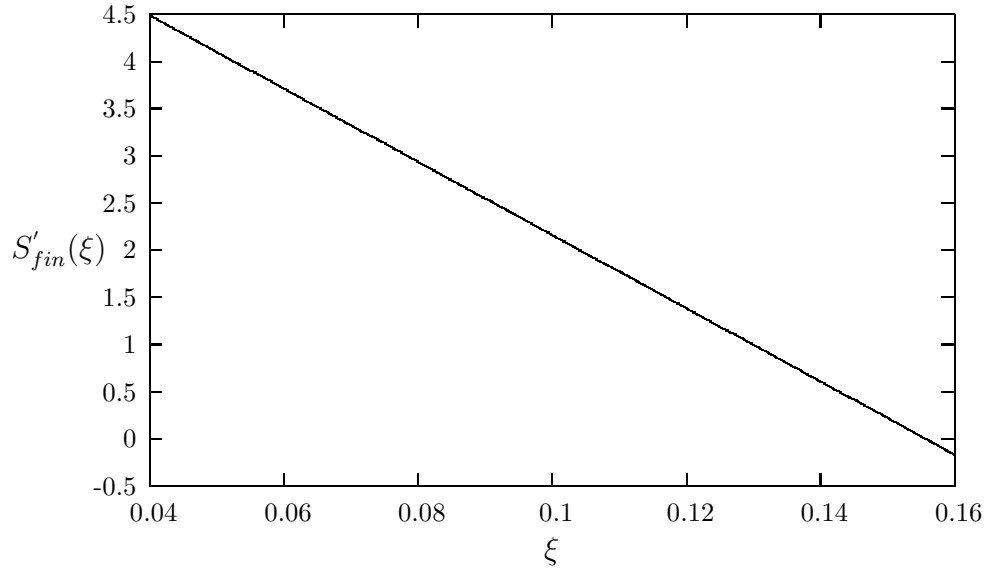


Figure 5: The Dependence $S'_{fin} = S'_{fin}(\xi)$, effectively of the Hubble parameter, at the end of this period.

These results imply that the presence of the conformal coupling ξ , for a positively curved manifold, act by slowing the inflationary evolution.

IV. Analytical Considerations

The system of equations (14) and (15) has three independent parameters. The first is δ which is related to the specific choice of the potential, i.e., the position of $V(0)$ relative to the Planck scale. The second is ξ which depends on the conformal coupling between the scalar curvature and the field. The third is $\bar{\lambda}$ which is a free parameter specifying the relation between the minimum scale factor with respect to the Planck length. Once one chooses values for these parameters, they specify a unique system for Eqs. (14) and (15). Then there exists a unique initial value for the field derivative $F'_0(\xi)$ that gives the solutions presented. This three-parameter freedom is equivalent to the choice of initial values for the scale factor, the Hubble parameter and the derivative of the scalar field, ϕ' at the end of this era, $(S_{fin}(\xi), S'_{fin}(\xi), F'_{fin}(\xi))$. The inflationary solution to be matched at this point is assumed to have $\phi_0 = 0$.

The behaviour of the system of Eqs. (14) and (15) can be approached analytically also. If we substitute into Eq. (14) the following ansatz (for $\bar{\lambda} = 1$)

$$F(T) = 1 - f_0 T, \quad f_0^2 = \frac{2}{3}, \quad (16)$$

$$S(T) = \sqrt{1 + \frac{4\xi}{f_0} T - 2\xi T^2}, \quad (17)$$

it is satisfied identically. When substituted into Eq. (15) it is an acceptable solution for small T . The parameter f_0 is ξ -invariant but this solution does not satisfy the requirement of zero Hubble parameter for $T = 0$. In this solution the scalar field decreases with time T , whereas the scale factor *increases* in a quasi-linear manner.

We present also a solution of Eqs. (7) and (10) that is not directly relevant to the case discussed here but it is interesting. We note that an exact solution to Eq. (7) is given by

$$R(t) = R_0 t + R_{min}, \quad \phi(t) = \frac{\phi_0 t}{(R_0 t + R_{min})^{3/2}}, \quad R_0^2 = 8\xi, \quad (18)$$

which shows that the presence of the conformal coupling ξ gives the non-zero scale factor evolution R_0 . This is important because it is in general not easy to find simple exact solutions of the Klein-Gordon equation, Eq. (7), for positively curved manifolds. The scalar field increases from zero value as well as the scale factor.

When Eq. (18) is substituted into Eq. (10) and one makes the approximation of small time ($t \ll 1$), the following relation should be satisfied

$$(1 - 6\xi)\phi_0^2 = 4R_{min}^3 V(0). \quad (19)$$

This yields a minimum value R_{min} which is positive for $\xi \in [0, \frac{1}{6}]$.

V. Discussion

We have considered the pre-inflation era of a positively curved FLRW Universe in the presence of the inflaton scalar field ϕ . This era begins from the Planck time up to

the time where the symmetry breaking phase transition takes place (e.g., the GUTs scale). The inflaton field is conformally coupled to the spacetime curvature scalar.

In inflationary theory the quantum fluctuations of scalar fields oscillate until their wavelength reaches H^{-1} . After that they freeze with amplitude $\delta\phi \sim \frac{H}{2\pi}$ [3,12]. So most of the wavelengths of the quantum fluctuations follow the evolution and the only assumption we have made is that their amplitude is suppressed as time evolves. In Eq. (7), expanding around a zero classical value for the scalar field, due to the linearity in ϕ , one sees that the quantum fluctuations obey the same equation. We have assumed that they obey also Eq. (10) at least in first order.

As regards the power spectrum of the primordial fluctuations, from Eq. (7) we see that expanding the scalar field in spherical harmonics, it should be possible, at least in principle, to detect the possible dependence on the conformal coupling parameter. This however would require a more detailed analysis in the framework of a positively curved manifold, which we defer for a future work.¹

The system has solutions with non-zero *minimum scale factor* and zero initial Hubble parameter. The quantum fluctuations of the inflaton field are assumed to be suppressed at the end of this era. This is of considerable importance because the minimum scale factor is several orders higher than the Planck length scale (in our case $R_{min} \simeq \delta^{-1}\chi \simeq 10^9 L_{Pl}$) and the Universe does not have to enter the Quantum Gravity era at all [9-11]. Also the finiteness of the scalar curvature invariant implies that one can continue backwards in time such solutions, particularly in the context of the so-called Pre-Big-Bang models [9].

As it was described in the previous section a particular inflationary solution (specifying the form of the scalar potential), with given initial data for the scale factor $R(t)$, the Hubble parameter $H(t)$, and the field time derivative $\phi'(t)$ (because we assume the initial value of the field to be taken as zero, $\phi(t) = 0$) can be matched to the end of this pre-inflationary evolution. Thus the whole solution is of the class $C^{(1)}$, i.e., continuously differentiable.

Finally it would be interesting to investigate the role of the possible conformal coupling between the scalar field and the spacetime curvature in more general inflationary models, that contain polynomials in the scalar curvature. Examples of these are the recently proposed models of [10-11]. Work along this line is in progress.

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